V. Shergotty

Basalt, 5 kg. seen to fall

Introduction

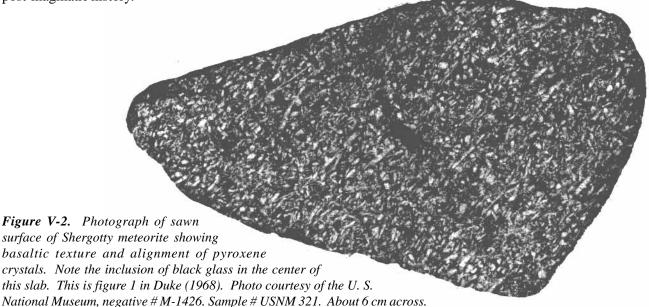
The Shergotty achondrite fell on August 25, 1865 at 9:00 a.m. near a town called Shergahti in Bihar State, India after detonations were heard (Graham *et al.*, 1985). Duke (1968) refers to several stones with fusion crusts. The main mass is at the Museum of the Geological Survey in Calcutta, India (figure V-1). In 1984, an international consortium was organized by J. C. Laul to study ~30 grams of Shergotty in detail (Laul, 1986 a,b).

The two shergottites (Shergotty, Zagami) are texturally and mineralogically similar to terrestrial diabases (although all of the plagioclase has been shocked to maskelynite), but quite distinct petrologically and chemically from the rest of the basaltic achondrites (Stolper *et al.* 1979). Stolper and McSween (1979) noted that Shergotty crystallized under relatively oxidizing conditions.

The Shergotty meteorite has been severely shocked. In fact, it is the type locality for maskelynite (dense plagioclase glass). It has proven to be very difficult to date the original crystallization event of Shergotty, perhaps because the sample has experienced a complex post-magmatic history.



Figure V-1. Photograph of main mass of Shergotty meteorite showing fusion crust and broken surface. Two saw cuts are visible. Sample is about 25 centimeters across. Photo kindly provided by Prof. N. Bhandari, Director, Physical Research Laboratory, Ahmedabad, India.



Petrography

The mineralogy and petrology of Shergotty have been comprehensively studied by many investigators (*e.g.* Tschermak, 1872; Binns, 1967; Duke, 1968; Smith and Hervig, 1979; Stolper and McSween, 1979; Nakamura *et al.*, 1982; Jagoutz and Wänke, 1986; Stöffler *et al.*, 1986; Lundberg *et al.*, 1988 and others).

Shergotty consists of dull gray-green pyroxene prisms up to 1 cm long with colorless lathlike and interstitial vitreous maskelynite which has replaced the feldspar (Duke, 1968). Shergotty exhibits a foliated texture produced by preferential orientation of pyroxene prisms and maskelynite grains (Duke, 1968, Stolper and McSween, 1979)(figure V-2). The texture was described as ophitic by Michel (1912), but the rather euhedral pyroxene does not enclose feldspar (Duke, 1968). The average grain size for Shergotty is only about 0.46 mm (figure V-3).

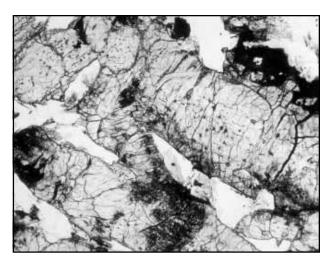


Figure V-3. Photomicrograph of thin section of Shergotty meteorite showing large clinopyroxene crystals separated by white maskelynite laths. Thin section kindly provided by Gordon McKay, JSC. Field of view is 2.2 mm.

Mineral Mode

	Stolper & McSween 1979		Duke 1968	Smith & Hervig 79	Stöffler et al 1986			
	mode	mode	norm	mode	mode			
Pyroxene	70.5	69.1	68.9	70	67			
Plagioclase	23.9	22.7	24.0	20	24			
Mesostasis	2.8	5.2		3	3			
Magnetite	2.0	2.5	2.3	2	2			
Ilmenite	0.5	0.3	1.8	0.5	0.2			
Pyrrhotite	0.3	0.2			0.4			
Phosphate	tr	tr	1.3	1	2			
silica					0.5			

In thin section, pyroxene prisms vary in color from light brown in the cores to yellow brown at the rims, reflecting a marked enrichment of iron at the rims (Duke, 1968). Maskelynite is transparent, colorless, conchoidally fractured, isotropic and is a pseudomorph after plagioclase (see below). Based on the texture, plagioclase crystallized late in the sequence of crystallization (Stolper and McSween, 1979; Treiman, 1985a).

Melt inclusions in pyroxene are an important assemblage in Shergotty (Nyquist *et al.*, 1979; Treiman, 1985a). They represent trapped magmatic liquid and may allow a determination of the age of entrapment during crystallization (see below).

The coexistence of fayalite, Ti-magnetite and a silica glass indicates late-stage crystallization under relatively high oxygen fugacity (Smith and Hervig, 1979).

Shergotty should be examined for the presence of shock-melted 'glass pockets' similar to those found in Zagami and EETA79001 (figure V-2).

Mineral Chemistry

Pyroxene: Pyroxene consists of separate grains of pigeonite and sub-calcic augite, both with a homogeneous interior and a strongly zoned Fe-rich rim (figure V-4). Smith and Hervig (1979), Stöffler *et al.* (1986) and Lundberg *et al.* (1988) have studied pyroxene zoning - both major and trace elements. There are two trends, Ca-rich from En₄₈Fs₁₉Wo₃₃ to En₂₅Fs₄₇Wo₂₈, and Ca-poor from En₆₁Fs₂₆Wo₁₃ to En₂₁Fs₆₁Wo₁₈. Mn increases along with Fe (figure V-5), and Smith and Hervig and Stolper and McSween (1979), found that the trend is towards an intercept of MnO near 0.2 % for low Fe. The preservation of extreme chemical zoning in pyroxenes indicates rapid cooling.

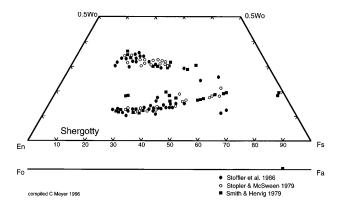


Figure V-4. Pyroxene composition diagram for Shergotty meteorite. Data combined from several sources including Stöffler et al. (1986), Stolper and McSween (1979) and Smith and Hervig (1979).

Müller (1993) studied "microstructures" in Shergotty clinopyroxene. Fine scale exsolution also argues for rapid crystallization in a lava flow or sill about 40-60 meters thick.

Pyroxferroite: Smith and Hervig (1979) reported minute grains of pyroxferroite included in maskelynite.

Maskelynite: Shergotty is the type specimen for maskelynite (Tschermak, 1872), a plagioclase glass formed by high shock pressures estimated at ~300 Kbars (Smith and Hervig, 1979) or 29 Gpa (Stöffler et al., 1986). Binns (1967) found that the maskelynite in Shergotty lacks any crystal structure, as evidenced by isotropism and absence of diffraction lines in long-exposure X-ray powder photographs. Duke (1968) noted that the refractive indices and specific gravities of maskelynite grains are variable, and suggested that individual crystals are zoned. Easton and Elliot (1977) reported the composition of "feldspar/glass" from Shergotty as An_{51,2}Ab_{46,5}Or_{2,3}. Stolper et al. (1979) determined that individual maskelynite grains are zoned from An₅₇Ab₄₂Or₁ to An₄₃Ab₅₃Or₄.

K-rich inclusions: The pyroxenes in Shergotty contain K (and Rb) rich inclusions (Nyquist *et al.*, 1979). It is these inclusions that apparently define the Rb/Sr isochron.

Phosphates: Most of the trace elements in the rock are in the phosphates. Fuchs (1969) gave an analysis of whitlockite in Shergotty. Jagoutz and Wänke (1986) give the composition of apatite and whitlockite. Smith and Hervig (1979) reported merrillite (?) in glass veins. Lundberg *et al.* (1988) analyzed both whitlockite and

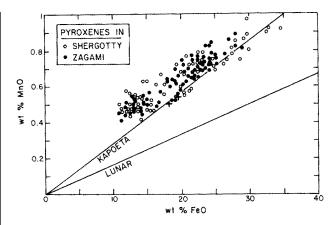


Figure V-5. Variation between MnO and FeO in pyroxenes in Shergotty and Zagami meteorites. This is figure 5 of Stolper and McSween (1979), GCA 43, 1480.

apatite for REE by ion microprobe. Wadhwa *et al.* (1994) compare whitlockites in the shergottites. In 1996, Dreibus and Wänke performed detailed leaching experiments on the phosphates.

Magnetite: Tschermak (1872) was the first to recognize primary magnetite in Shergotty. This important observation, overlooked for many years, indicates a high degree of oxidation at the time of crystallization.

Baddeleyite: Smith and Hervig (1979) reported rare, small (~1 micron) grains of baddeleyite. Lundberg *et al.* (1988) found that baddeleyite had a flat REE pattern.

Hercynite spinel: Treiman (1985) reported Fe, Alrich spinel in magmatic inclusions in pyroxene.

Amphibole: Treiman (1984, 1985) reported pleochroic amphibole in the magmatic inclusions in pyroxene. Amphibole has been analyzed for D/H ratio by Watson *et al.* (1994).

Sulfide: The sulfide in Shergotty was recognized to be pyrrhotite by Tschermak (1872). Smith and Hervig (1979) describe it as "homogeneous with molar Fe:S of 0.94, and 0.12 wt % Ni, 0.03 Cu." It is found in the rims of the pyroxene and in the mesostasis.

Silica: Duke (1968) reported silica (cristobalite ?), but could not find x-ray lines for coesite. Stöffler *et al.* (1986) reported a large grain (~200 microns) of α-quartz adjacent to maskelynite and one grain of possible stishovite (?).

Fayalite: Smith and Hervig (1979) found that fayalite (Fo₁₀) was present in the mesostasis adjacent to magnetite-ilmenite grains.

Whole-rock Composition

The major element composition of Shergotty was determined by Tschermak (1872), Duke (1968), McCarthy *et al.* (1974), and Dreibus *et al.* (1982). The Shergotty Consortium (Laul *et al.*, 1986) obtained additional chemical data on Shergotty (table V-1). The REE data are compared in figure V-6. Laul *et al.* (1986) noted a slight difference in the REE abundance from fragments B and C from opposite sides of the main mass, possibly indicating some heterogeneity in the sample.

In addition to the data table, Allen and Clark (1977) determined 58 ppm F and reported 56.2 ppm F by Reed and Jovanovic (1969). Gibson and Moore (1983) and Gibson *et al.* (1985) reported two values for S in Shergotty (360 and 2170 ppm respectively). Dreibus *et al.* (1983) and Laul *et al.* (1986) reported 1330 and 1579 ppm S. Burgess *et al.* (1989) found 1930 ppm S and suggested that the wide range of reported values is due to sample heterogeneity (see figure II-7). Dreibus and Wänke (1987) reported 36 ppb I, 890 ppb Br, and 108 ppm Cl. Jovanovic and Reed (1987) reported 470 ppb Hg (probably Museum contamination).

Yang and Epstein (1985) reported 180 ppm H₂O in Shergotty. Karlsson *et al.* (1992) reported 640 ppm H₂O, but recognized that some of this may be terrestrial (*see section on Other Isotopes*). Gooding *et al.* (1990) determined the thermal release pattern for several volatile species.

Stolper (1979) noted the remarkable similarity of Shergotty and Zagami with terrestrial basalts for a wide range of elements (figure V-7). However, Dreibus *et al.* (1992) noted that Shergotty had much less Ir than the other Martian meteorites and/or terrestrial "upper mantle."

Note: The values for the composition of Shergotty reported by Schmitt et al. (1972) are so far off that Smith and Steele (1979) asked if they analyzed the wrong sample!! Also the values given for Shergotty in the book on Basaltic Volcanism (Lofgren and others, 1981) do not seem to be consistent with literature values.

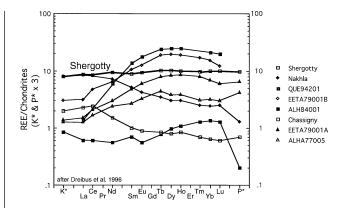


Figure V-6. REE diagram comparing Shergotty with other Martian meteorites. This diagram is redrafted from Dreibus et al. (1996).

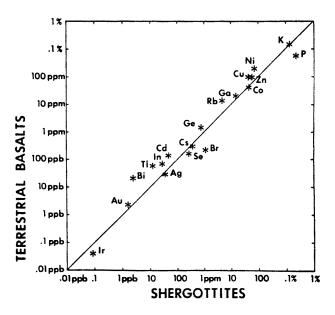


Figure V-7. Comparison of the composition of Shergotty with average composition of terrestrial basalts by Stolper (1979), EPSL 42, 239.

Radiogenic Isotopes

The original crystallization age of Shergotty has proven to be difficult to determine. Geiss and Hess (1958) originally determined a K/Ar age of 560 Ma (+60-150), and suggested that the Rb/Sr technique be used (page 233). Twenty one years later, Nyquist *et al.* (1979) *finally* determined an internal Rb/Sr isochron for Shergotty of 165 ± 11 Ma with high 87 Sr/ 86 Sr = 0.72260 ± 0.00012 (figure V-8) ($\lambda_{RB} = 1.39 \times 10^{-11}$ year⁻¹). However, Bogard *et al.* (1979) attempted to determine the 39 Ar/ 40 Ar age and found that the maskelynite gave a plateau age of 254 ± 10 Ma (figure V-9).

Table V-1. Chemical analyses of Shergotty.

	Treiman 86	Dreibus 82	McCarthy 74	Jerome 70	Philpotts70	Tscher. 1872	Duke 1968	Smith 84	Shih 82	Ma 81	Warren 87
weight	0.1-0.2 g		1 g	1.065 g						(Irving)	
SiO2 %		51.36	50.36 (f)	51.48 (a)		50.21 (e)	50.1 (e)				
TiO2 Al2O3		0.87 7.06	0.85 (f) 7.03 (f)	0.81 (a) 5.75 (a)		5.9 (e)	0.92 (e) 6.68 (e)				6.42
Fe2O3							1.49 (e)				
FeO MnO		19.41 0.52	19.34 (f) 0.54 (f)	19.8 (a) 0.49 (a)		21.85 (e)	18.66 (e)				18.93 0.515
CaO		10	9.58 (f)	10.25 (a)		10.41 (e)	0.5 (e) 10.03 (e)				8.81
MgO		9.28	9.27 (f)	9.35 (a)		10 (e)	9.4 (e)				9.78
Na2O		1.29	n.d.	1.31 (a)	0.10 (1)	1.28 (e)	1.28 (e)				1.34
K2O P2O3		0.164 0.8	0.18 (f) 0.47 (f)	0.13 (a)	0.18 (d)	0.51 (e)	0.16 (e) 0.71 (e)				0.178
sum		100.754	97.62	99.37		100.16	99.93				
Li ppm C		5.6 620							4.1 (d)		
F		41.6									
S		1330									
Cl So		108		40 (h)	Jerome 70		72 (b)	50		5 2 (a)	52
Sc V		58.9		48 (b) 380 (b)	50.3 (c)		72 (b) 340 (b)	52		53 (c)	290
Cr		1389	2100 (f)	1500 (b)	1400 (c)	Laul 72	1150 (b)				1330
Co Ni	55 (g)	39 83		37 (b) 81 (b)	39.5 (c)	63 (g)	42 (b) 100 (b)	38 (g) 56		35 (c) 56 (c)	39 80
Cu	55 (g)	26		11 (b)		11.9 (g)	100 (0)	54		50 (C)	30
Zn	68.3 (g)	83				76 (g)		76 (g)			54
Ga Ge	0.69 (g)	14.7				13.9 (g)		15 (g)			16.7
As	0.07 (g)	0.025						6.03 (g)			
Se	0.37 (g)	0.41				0.242 (g)		0.47 (g)			
Br Rb	1.06 (g) 9.5 (g)	0.891 (j) 6.84			0.13 (d)	1.07 (g) 6.1 (g)		5.7	6.22 (d)		
Sr).5 (g)	51		35 (b)	50.7 (d)	0.1 (g)	8 (b)	5.7	51 (d)		
Y				<10 (b)			19 (b)				
Zr Nb							41 (b) 22 (b)				
Мо							22 (0)				
Pd ppb	1.7 (g)					262 ()		110 ()			
Ag ppb Cd ppb	16.9 (g) 12.9 (g)					263 (g) 44 (g)		110 (g) 340 (g)			
In ppb	26.5 (g)					31.3 (g)		23 (g)			
Sb ppb	1.7 (g)	<5						27 (g)			
Te ppb I ppm	2.5 (g)	0.036 (j)						3.2 (g)			
Cs ppm	0.529 (g)	0.405			Schnetz. 69	0.38 (g)		0.44 (g)			
Ba		35.5		40 (b)	32 (d)		6 (b)		29.4 (d)	2.10 (-)	1.00
La Ce		2.29 5.54		1.93 (c)	5.89 (d)				1.5 (d) 3.51 (d)	2.18 (c)	1.77
Pr		0.86									
Nd Sm		4.5 1.37		1.37 (c)	4.96 (d) 1.89 (d)				2.6 (d)	1.36 (c)	1 27
Eu		0.564		0.55 (c)	0.643 (d)				0.43 (d)	0.53 (c)	
Gd		2.2			2.8 (d)				1.64 (d)		
Tb Dy		0.44 2.94			3.38 (d)				2.16 (d)	0.36 (c)	
Но		0.56									
Er Tm		1.87 0.38			1.89 (d)				1.33 (d)		
Yb		1.69		1.4 (c)	1.8 (d)				1.19 (d)	1.59 (c)	
Lu Hf		0.25 1.97		0.34 (c) 2.15 (c)				2	0.176 (d)	0.262 (c) 2 (c)	0.23
Ta		0.25		· - (-)				0.27		0.27 (c)	
W ppb	0.044 (%)	480									
Re ppb Os ppb	0.044 (g) <.0023 (g)										
Ir ppb	0.0285 (g)	0.4 (k)				1.55 (g)					0.07
Au ppb Tl ppb	5 (g)	16				88 (g)		12 (g)			0.97
Bi ppb	12.9 (g) 0.7 (g)	12 (1) 2 (1)	Morgan 73	Chen 86	Chen 86	13.4 (g) 3.7 (g)		11 (g) 1.6 (g)			
Th ppm		0.39	0.465 (e)	0.439 (d)	0.398 (d)					0.35 (c)	
U ppm	0.129 (g)	0.116	0.112 (e)	0.12 (d)	0.099 (d)			0.171 (g)			

technique: (a) semi-micro wet chem., (b) emission spec., (c) INAA, (d) isotope dilution mass spec., (e) wet chem., (f) XRF, (g) RNAA (h) INAA of less than 44 micron fraction, (i) mistake(?) see note in text, (j) Dreibus et al 1985, (k) Burghele 1983, (l) from McSween 1985, (m) elec. probe, fused sample

Table V-1. Chemical analyses of Shergotty continued.

	Schmitt 72	Schmitt 72	Schmitt 72	Laul 86	Laul 86	Laul 86	Laul 86	Laul 86	Laul 86	Stolper 79	Jagoutz86	Jagoutz86
weight	47 mg	915 mg	1.457 g	72 mg	90 mg	139.8 mg	130.7 mg	94 mg	240 mg			
SiO2 %	(i)	(i)	(i)			49.5 (f)				50.4 (m)		
TiO2 Al2O3				0.8 (c) 7.14 (c)	0.7 (c) 6.8 (c)	0.87 (f) 7.59 (f)				0.81 (m) 6.89 (m)		
Fe2O3	1501 ()	1415 ()					20.1 ()					
FeO MnO	16.21 (c) 0.53 (c)	14.15 (c) 0.2 (c)		19.6 (c) 0.52 (c)	19.9 (c) 0.535 (c)	19.8 (c) 0.518 (c)	20.1 (c) 0.54 (c)			19.1 (m) 0.5 (m)		
CaO				9.52 (c)	9.72 (c)	9.63 (f)	9.88 (c)			10.1 (m)		
MgO Na2O	1.71 (c)	1.81 (c)	1.25 (c)	9.5 (c) 1.4 (c)	9 (c) 1.3 (c)	8.95 (f) 1.47 (c)	1.28 (c)			9.27 (m) 1.37 (m)		
K2O			(4)	0.19 (c)	0.15 (c)	0.189 (c)	0.165 (c)			0.16 (m)		
P2O3 sum						0.72 (f) 98.52				98.6		
Li ppm						3.3 (c)						
C F						430						
S						1570						
Cl Sc	49 (c)	35 (c)	26 (c)	54 (c)	56.5 (c)	108 (c) 53.8 (c)	57.8 (c)					
V	1) (0)	33 (0)	20 (0)	260 (c)	265 (c)		37.0 (c)					
Cr Co	1150 (c)	350 (c) 62 (c)	960 (c)	2616 (c) 37.5 (c)	2821 (c) 38.5 (c)	2923 (c) 37.9 (c)	2675 (c) 39.5 (c)	37.2 (g)	37.2 (g)			
Ni	1150 (c)	02 (0)	700 (C)	70 (c)	70 (c)	81 (c)	88 (c)	37.2 (g)	31.2 (g)			
Cu Zn		139 (c)		67 (c)	68 (c)	26 (c)		69.8 (g)	68.4 (g)			
Ga				07 (C)	06 (C)	17.6 (c)	16.5 (c)	69.8 (g) 17.5 (g)	68.4 (g) 15.7 (g)			
Ge						0.025 (a)						
As Se						0.025 (c)		0.38 (g)	0.4 (g)			
Br						0.89 (c)	0.8 (c)				6.61 (1)	7.52 (1)
Rb Sr				54 (c)	45 (c)			7.27 (g)	6.11 (g)		6.61 (d) 60.12 (d)	7.53 (d) 50.21(d)
Y												
Zr Nb				67 (c)	50 (c)							
Mo						0.37 (c)						
Pd ppb Ag ppb								10.5 (g)	6.8 (g)			
Cd ppb								44 (g)	14 (g)			
In ppb Sb ppb								25 (g) <20 (g)	23 (g) 8.9 (g)			
Te ppb								19 (g)	4.3 (g)			
I ppm Cs ppm						0.036 (c) 0.48 (c)	0.42 (c)	0.46 (g)	0.4 (g)		0.53 (d)	0.56 (d)
Ba				32 (c)	27 (c)	40 (c)	39 (c)	0.10 (g)	o (g)		0.00 (a)	0.00 (4)
La Ce				2.44 (c) 5.8 (c)	1.95 (c) 4.7 (c)	2.44 (c) 6.4 (c)	1.97 (c) 5.6 (c)					
Pr				0.88 (c)	0.7 (c)	0.1 (c)	3.0 (c)					
Nd Sm				4.2 (c) 1.6 (c)	3.3 (c) 1.3 (c)	4.7 (c) 1.66 (c)	4.3 (c) 1.46 (c)				4.18 (d) 1.63 (d)	3.45 (d) 1.4 (d)
Eu				0.65 (c)	0.55 (c)	0.65 (c)	0.55 (c)				1.05 (u)	1.4 (u)
Gd Tb				2.6 (c) 0.52 (c)	2.1 (c) 0.41 (c)	2.8 (c) 0.52 (c)	2.5 (c) 0.42 (c)					
Dy				3.7 (c)	3 (c)	4.8 (c)	3.1 (c)					
Ho Er				0.73 (c)	0.6 (c)	0.86 (c)	0.8 (c)					
Tm				0.3 (c)	0.25 (c)	0.31 (c)	0.3 (c)					
Yb				1.8 (c) 0.26 (c)	1.5 (c)	1.79 (c)	1.62 (c)					
Lu Hf				0.26 (c) 2.15 (c)	0.22 (c) 1.5 (c)	0.26 (c) 2.23 (c)	0.24 (c) 1.83 (c)					
Ta				0.29 (c)	0.18 (c)	0.27 (c)	0.23 (c)					
W ppb Re ppb						500 (c)	400 (c)					
Os ppb				.e / \	.e / \	2 (2 ()					
Ir ppb Au ppb				<5 (c) <6 (c)	<5 (c) <6 (c)	<3 (c) 0.9 (c)	<3 (c) 6 (c)	5.4 (g)	0.81 (g)			
Tl ppb				(-)	(-)	(-)	(-)	14 (g)	0.15 (g)			
Bi ppb Th ppm				0.37 (c)	0.25 (c)	0.37 (c)	0.29 (c)	1.2 (g)	0.47 (g)			
U ppm				0.095 (c)	0.095 (c)	0.095 (c)	0.095 (c)	0.13 (g)	0.055 (g)			

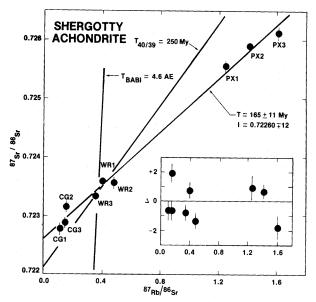


Figure V-8. Rb-Sr isochron diagram for mineral separates and whole rock samples from Shergotty meteorite. This is figure 1 in Nyquist et al. (1979), GCA 43, 1059.

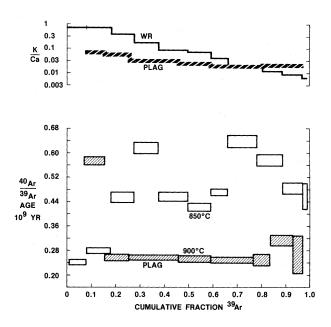


Figure V-9. Ar plateau diagram for the age of Shergotty meteorite. This is figure 1 in Bogard et al. (1979), GCA 43, 1049.

Jagoutz and Wänke (1986) repeated the work of Nyquist *et al.* and reported a Rb/Sr age of 167 Ma using maskelynite and mesostasis with high I⁸⁷Sr/⁸⁶Sr = 0.72263 ± 0.00005 . They also reported two Sm-Nd ages (147 \pm 20 Ma for leaches of the "whole rock" and 360 \pm 16 Ma for the pyroxene separates and leaches)(figure V-10). Shih *et al.* (1982) were unable to use the Sm-Nd system to determine a precise crystallization age for Shergotty.

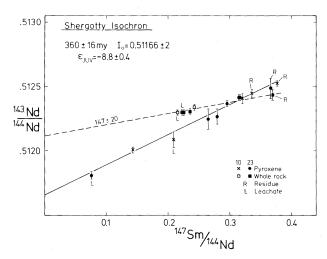


Figure V-10. Sm-Nd isochron diagram for pyroxene separates and their leachates from Shergotty meteorite. This is figure 6 in Jagoutz and Wänke (1986), GCA 50, 946.

By leaching "whole rock" samples of Shergotty, Chen and Wasserburg (1986a) obtained a U-Pb "isochron" of 200 ± 4 Ma (figure V-11) and a Th-Pb "isochron" of 205 ± 8 Ma (figure V-12). These leach experiments probably attacked the phosphates in the sample (see section on Experiments below). By leaching the more refractory "pyroxene" separates, Chen and Wasserburg obtained 600 ± 20 Ma for U-Pb and 437 ± 36 Ma for Th-Pb.

There are many more interpretations of the age data for Shergotty than can be summarized here (see Jones (1986,1989) for an interesting approach). Note that by careful petrography Müller (1993) concluded that "the event around 160 Ma is not responsible for the shock metamorphism displayed by Shergotty, but may be the crystallization age."

Cosmogenic Isotopes and Exposure Ages

Heymann *et al.* (1968) reported an exposure age of ~2 Ma and Bogard *et al.* (1984b) calculated an exposure age of ~2.6 Ma for Shergotty. Nishiizumi *et al.* (1986b) reported an exposure age of 2.2 Ma. Pal *et al.* (1986) determined an exposure age of 2.7 ± 0.9 Ma using ¹⁰Be. From cosmic-ray produced ³He, ²¹Ne and ³⁸Ar, Eugster *et al.* (1996) calculated an exposure age for Shergotty of 2.8 Ma (2.5 Ma using ³He only) and concluded that Shergotty was "*ejected from Mars simultaneously with the other basaltic shergottites QUE94201 and Zagami*" about 2.6 Ma ago.

Other Isotopes

Taylor et al. (1965) found that the oxygen isotopes

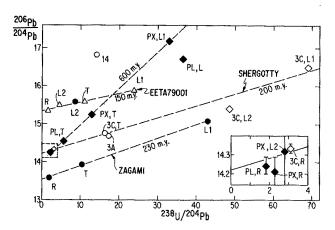


Figure V-11. U-Pb isochron diagram for Shergotty meteorite from Chen and Wasserburg (1986c), GCA 50, 959.

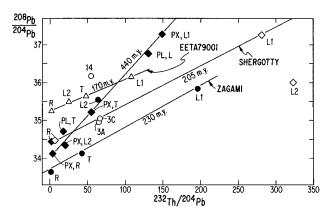


Figure V-12. U-Th-Pb diagram for Shergotty meteorite from Chen and Wasserburg (1986c), GCA 50, 959.

¹⁸O/¹⁶O in Shergotty were different from howardites and eucrites. Clayton and Mayeda (1983, 1996) reported the oxygen isotopes for Shergotty and revised the data of Clayton *et al.* (1976). Clayton and Mayeda (1986) reported additional analyses, including mineral separates. Clayton (1993) reports the ¹⁸O/¹⁶O composition of plagioclase and pyroxene from Shergotty and calculated the equilibrium temperature.

Watson *et al.* (1994) reported the deuterium contents of hydrous amphiboles in Shergotty. Leshin *et al.* (1996) found water released from Shergotty had high D/H ratios (figure V-13).

Fallick *et al.* (1983) reported $\delta^{15}N$ as low as -40 % for some steps of combustion. However, Becker and Pepin (1986) found that nitrogen was within a few per mil of terrestrial when effects of contamination and spallation are considered. Molini-Velsko *et al.* (1986) reported the isotopic composition of Si and found it normal.

Chen and Wasserburg (1986) reported the Pb isotopes in Shergotty and concluded that the parent body (Mars) was enriched in ²⁰⁴Pb and other volatiles.

Lugmair *et al.* (1996) studied the 53 Mn - 53 Cr systematics in a bulk sample of Shergotty and found that it gave the same 53 Cr/ 52 Cr excess of $0.23 \pm 0.10\epsilon$ as similar data from ALH84001 (figure V-14).

Experiments

Dreibus and Wänke (1996) have reported interesting experiments on the leaching of Shergotty with various solutions, showing that the phosphates readily dissolve, releasing many trace elements to the leach (see section on Radiogenic Isotopes above). This is applicable to what may have occurred on the surface of Mars, and may explain discordance in radiometric age determinations.

Stolper and McSween (1979) experimentally determined the phase equilibria and crystallization sequence for shergottites.

Thellier's method of paleointensity determination was applied to Shergotty (Cisowski, 1986), but the result was inconclusive, probably because of the high degree of shock suffered by Shergotty.

Fission-track mapping by Lundberg *et al.* (1988) found that only ~20 % of the U was located in the phosphates.

Bunch and Cohen (1968), Sippel (1971) and Sears and Hasan (1985) have studied the thermoluminescence of maskelynite (figure V-15).

Shock Effects

The shock history of Shergotty has been much discussed. Duke (1968) showed that Shergotty could not have been heated above 400° C following the shock. Lambert and Grieve (1984) and Gibbons and Ahrens (1977) concluded that the high index of refraction of maskelynite in Shergotty was the result of two shock events, one that originally produced the maskelynite and a second lower pressure shock thought responsible for the index of refraction. However, Stöffler *et al.* (1986) concluded that only one shock is needed and estimated the shock pressure was 29 ± 1 GPa with a post shock temperature of $200 \pm 20^{\circ}$ C (figure V-16). Müller (1993) concluded that the magnitude of the shock event was not sufficient to reset the Rb-Sr age.

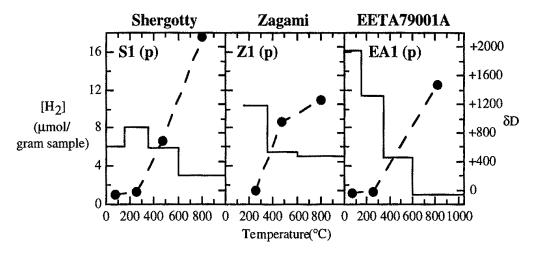


Figure V-13. Hydrogen isotope composition of water released from Shergotty, Zagami and EETA79001 meteorites. This is a copy of figure 3 in Leshin et al. (1996), GCA **60**, 2641.

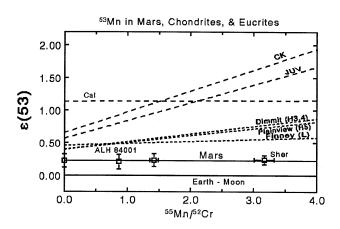


Figure V-14. Isotopic composition of Cr for Earth, Moon, Martian and other classes of meteorites including Shergotty. This is a copy of a figure in Lugmair et al. (1996), LPS XXVII page 786.

Hörz *et al.* (1986) found that Shergotty pyroxenes showed continuous diffraction rings, line broadening and other evidence of substantial lattice disorder. Stöffler *et al.* reported one grain of silica that might have been stishovite (?).

Processing

The main mass of Shergotty (~3.6 kg.) is curated by the Geological Survey of India, in Calcutta. About 60 % of the surface of this large piece is covered with fusion crust (figure V-1). Over the years, there have been two saw cuts and several large pieces broken off. Duke (1968) refers to more than one piece of Shergotty.

In 1986, J. C. Laul and his colleagues reported the results of the international consortium study on Shergotty in *Geochimica et. Cosmochimica Acta*, vol.

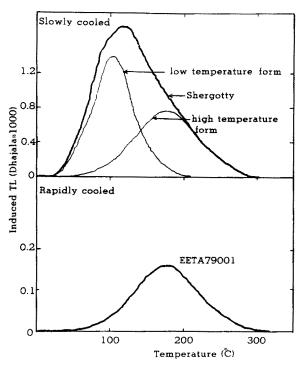


Figure V-15. Experimental thermoluminescence glow curves for Shergotty and EETA79001 meteorites. This is figure 7 from Hasan et al. (1985), GCA 50, 1037.

50. The samples for the Shergotty Consortium (~30 g) were processed in the clean labs at JSC from three chips (labeled A, B and C) provided by the Geological Survey of India. Fragment A (5 g) and fragment C (13 g) were chipped from the main mass next to each other, while fragment B (12 g) was obtained from one of the areas sawn off the other side of the main mass. Figure V-17 shows the location of three chips from the main mass. Powdered samples were prepared from chips B and C for chemical analysis by the consortium (Laul, et al. 1986). Figure V-18 attempts to show the recent history of the various pieces of the sample.

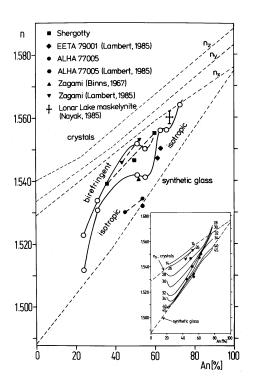


Figure V-16. Refractive index measurements on plagioclase that has been shocked experimentally compared to Martian meteorites. This is figure 12 from Stöffler et al. (1986), GCA **50**, 897.

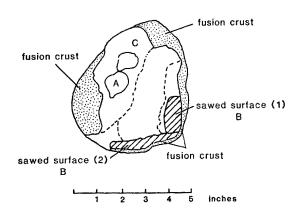


Figure V-17. Sketch of main mass of Shergotty showing the location of samples studied by Shergotty Consortium. This is figure 5 in Laul (1986), GCA 50, 878.

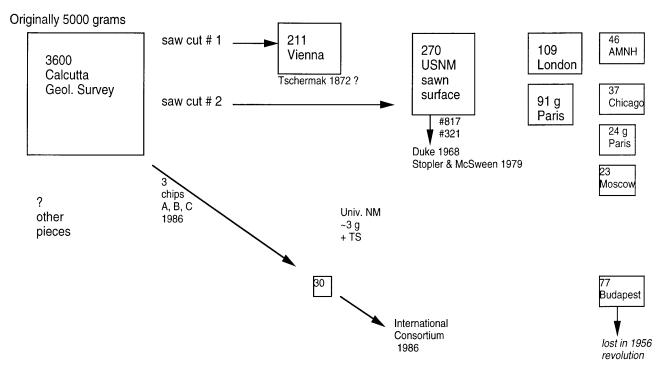


Figure V-18. Distribution of pieces of Shergotty meteorite (preliminary diagram). Was there more than one piece originally? Which research has been done on which piece?